

A spectroscopic investigation of early-type stars in the young open cluster Westerlund 2[★]

G. Rauw^{1,★★}, H. Sana², and Y. Nazé^{1,***}

¹ Groupe d'Astrophysique des Hautes Energies, Institut d'Astrophysique et de Géophysique, Université de Liège, Allée du 6 Août, Bât B5c, 4000 Liège, Belgium

² Sterrenkundig Instituut 'Anton Pannekoek', Universiteit van Amsterdam, Science Park 904, 1098 XH Amsterdam, The Netherlands

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ABSTRACT

Context. The distance of the very young open cluster Westerlund 2, which contains the very massive binary system WR 20a and is likely associated with a TeV source, has been the subject of much debate.

Aims. We attempt a joint analysis of spectroscopic and photometric data of eclipsing binaries in the cluster to constrain its distance.

Methods. A sample of 15 stars, including three eclipsing binaries (MSP 44, MSP 96, and MSP 223) was monitored with the FLAMES multi-object spectrograph. The spectroscopic data are analysed together with existing *B V* photometry.

Results. The analysis of the three eclipsing binaries clearly supports the larger values of the distance, around 8 kpc, and rules out values of about 2.4 – 2.8 kpc that have been suggested in the literature. Furthermore, our spectroscopic monitoring reveals no clear signature of binarity with periods shorter than 50 days in either the WN6ha star WR 20b, the early O-type stars MSP 18, MSP 171, MSP 182, MSP 183, MSP 199, and MSP 203, or three previously unknown mid O-type stars. The only newly identified candidate binary system is MSP 167. The absence of a binary signature is especially surprising for WR 20b and MSP 18, which were previously found to be bright X-ray sources.

Conclusions. The distance of Westerlund 2 is confirmed to be around 8 kpc as previously suggested based on the spectrophotometry of its population of O-type stars and the analysis of the light curve of WR 20a. Our results suggest that short-period binary systems are not likely to be common, at least not among the population of O-type stars in the cluster.

Key words. Open clusters and associations: individual: Westerlund 2 – Stars: early-type – binaries: eclipsing – Stars: fundamental parameters – Stars: individual: WR 20b – Stars: individual: Cl* Westerlund 2 MSP 18

1. Introduction

The open cluster Westerlund 2 lies in a blowout region of the giant H II complex RCW 49. Interest in the stellar population of Westerlund 2 was triggered by two independent observational studies. On the one hand, RCW 49 was observed with the Infrared Array Camera aboard *Spitzer* in the framework of the GLIMPSE legacy survey (Churchwell et al. 2004). These observations revealed strong evidence of ongoing star formation activity and underlined the need for a detailed study of the properties of the early-type stars in the cluster core and their impact on the surrounding nebula. On the other hand, WR 20a - one of the two Wolf-Rayet stars in RCW 49 - was found to be a very massive eclipsing binary consisting of two WN6ha stars with individual masses of about 80 M_{\odot} (Rauw et al. 2004, 2005; Bonanos et al. 2004). Since previous spectroscopic studies revealed only one O6: and six O7: stars, in the cluster core (Moffat, Shara & Potter 1991), the existence of a pair of stars that massive was somewhat unexpected.

The nature of Westerlund 2 thus remains to be established: is it a super-cluster maybe harbouring dozens of very massive early O-type stars, or is it "just" a medium-size agglomerate

that produced the very massive WN6ha + WN6ha binary system WR 20a more or less by chance? ¹ This question called for a re-investigation of the population of early-type stars in this cluster. A first step in this direction was performed by Rauw et al. (2007, hereafter Paper I), who presented a photometric monitoring of the cluster as well as a spectroscopic snapshot study of its brightest members. This study allowed us to discover three previously unknown eclipsing binaries among the likely cluster members and to derive spectral types between O3 and O6.5 (i.e. significantly earlier than previously thought) for the twelve brightest O-stars in the cluster. Moreover, we inferred a spectrophotometric distance of 8.0 ± 1.4 kpc. While this distance measurement is in excellent agreement with that of WR 20a (Rauw et al. 2005), it is nevertheless problematic in several ways. Indeed, despite the much earlier spectral types inferred for the most massive cluster members, the current census of early-type stars in Westerlund 2 accounts for only 20% of the ionizing photons required to explain the observed radio emission of the RCW 49 complex. Moreover, a near-IR photometric study of the cluster by Ascenso et al. (2007) resulted in a distance estimate of 2.8 kpc, inferred from a comparison between their *JHK_s* data with pre-main sequence evolutionary tracks. The origin of these discrepancies remains unclear. More recently, yet another distance determination was obtained. From the mean velocity and

Send offprint requests to: G. Rauw

[★] Based on observations collected at the European Southern Observatory (Cerro Paranal, Chile)

^{★★} Honorary Research Associate FRS-FNRS (Belgium)

^{***} Research Associate FRS-FNRS (Belgium)

Correspondence to: rauw@astro.ulg.ac.be

¹ Indeed, the existence of a relation between the mass of an open cluster and that of its most massive member remains disputed (see Weidner, Kroupa & Bonnell 2010 and references therein).

velocity spread of molecular gas in the molecular cloud associated with RCW 49, Furukawa et al. (2009) inferred a kinematic distance of $5.4^{+1.1}_{-1.4}$ kpc. These authors further suggested that the collision between two molecular clouds might have triggered the formation of the Westerlund 2 cluster.

Determining the correct distance of this cluster is a fundamental issue: if the distance were confirmed to be 8.0 kpc, Westerlund 2 might be one of the most massive clusters of our Galaxy (along with NGC 3603, Cyg OB2, and Westerlund 1). Otherwise, if the distance were established to be much shorter than 8.0 kpc, WR 20a would be left in relative isolation, thereby challenging the theories of competitive accretion (Bonnell, Vine & Bate 2004) and the link between the mass of a cluster and the mass of its most massive member (Weidner et al. 2010). The distance is also of fundamental importance to understand the link with the TeV source HESS J1023–575 (Aharonian et al. 2007), which is spatially consistent with the bright *Fermi*-LAT pulsar PSR J1023.0–5746 (Abdo et al. 2010, Ackermann et al. 2011). At much longer wavelengths, a peculiar jet as well as arc-like molecular (^{12}CO) emission feature were reported by Fukui et al. (2009). The arc is well centred on the TeV source HESS J1023–575, and Fukui et al. (2009) argue that both the TeV and GeV γ -ray sources and the molecular features are likely to be associated with the same high-energy event. These authors accordingly suggest that these features could be due to an anisotropic supernova explosion of one of the most massive members of Westerlund 2 or to remnant objects such as a pulsar wind nebula or a microquasar. On the basis of the detection of significant emission during the off-pulse intervals with *Fermi*-LAT, Ackermann et al. (2011) argued in favour of the pulsar wind nebula scenario. In either case, the association of such a remnant with Westerlund 2 would imply that the most massive members of the cluster have already gone supernova, assuming star formation in this region was coeval. It must be stressed however that the preferred distance of PSR J1023.0–5746 is 2.4 kpc (Ackermann et al. 2011). Therefore, if the distance of Westerlund 2 were established to be much larger than 2.5 kpc, this would cast serious doubt on a connection between the pulsar and the cluster.

In this paper, we present the results of a spectroscopic follow-up campaign of the eclipsing binaries and the stars in the core of Westerlund 2. This campaign was specifically designed to address the questions outlined above. Throughout this work, we use the naming convention introduced by Moffat et al. (1991), except for objects that are discussed here for the first time. This paper is organized as follows. In Sect. 2, we provide a brief overview of our observing campaign and the data reduction. Section 3 presents our results concerning spectral classification, variability, and multiplicity for each star. The implications of these results are finally discussed in Sect. 4.

2. Observations

We obtained eight observations of the newly discovered eclipsing binaries in Westerlund 2 (Paper I), as well as several O-star binary candidates with the FLAMES multi-object spectrograph on the Very Large Telescope (VLT) UT2 (see Table 1). The observations were performed in service mode over an interval of one week in March 2008. Each observation lasted for an integration time of 2775 s and the observing conditions were good with a seeing of better than 1.1 arcsec. FLAMES offers the possibility to simultaneously acquire high-resolution spectra of up to eight targets with the UVES spectrograph and medium-

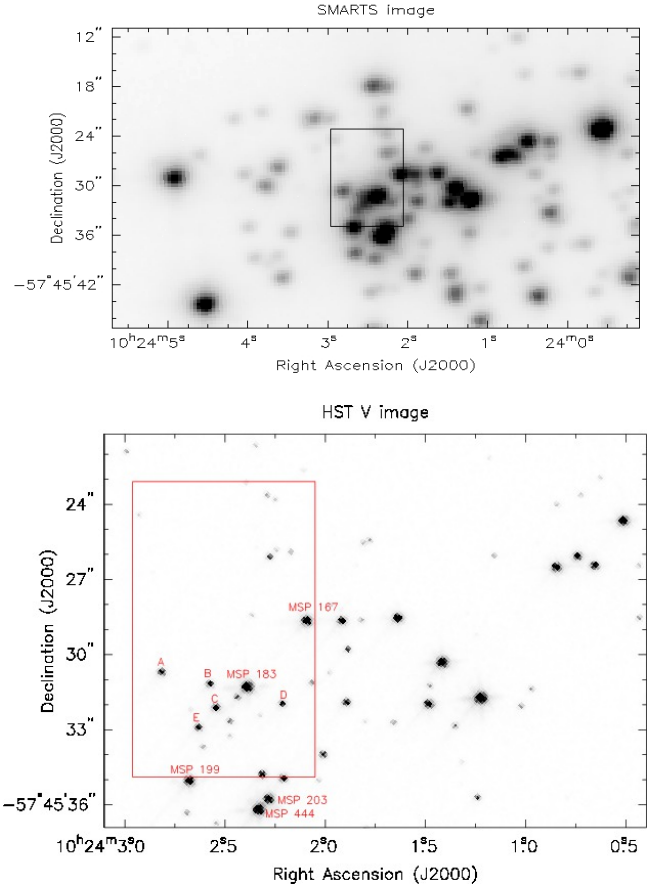


Fig. 1. Illustration of the position of the ARGUS unit (box) on the core of Westerlund 2. The top image uses our combined *B* and *V* image from Paper I (obtained with the SMARTS 1.3 m telescope at CTIO), whilst the bottom image uses an archive WFPC2 HST (F555W) image. The various objects analysed in the ARGUS data are labeled in the bottom figure.

resolution spectra of some additional targets with the GIRAFFE spectrograph.

Our campaign was divided into two parts (run A and B) with two different instrumental set-ups for the GIRAFFE spectrograph. During all our observations, five well-isolated targets (WR 20b, MSP 18, as well as the three eclipsing systems MSP 44, 96, and 223) were observed with FLAMES-UVES fibres. The UVES echelle spectrograph was operated in the standard setup Red 580, providing coverage of the wavelength domains 4842 – 5728 Å and 5843 – 6800 Å with a resolving power of 47000. During run A (five observations), the GIRAFFE spectrograph was fed by three integral field units (IFU), allowing us to monitor three additional O-type stars (MSP 171, 182, and 203). Whilst MSP 182 and 171 are well-isolated targets, MSP 203 lies in a more crowded region near the cluster core and its spectrum is heavily blended with that of MSP 444 (see Moffat et al. 1991). During run B (three observations), GIRAFFE was used with the ARGUS system. ARGUS consists of an array of 14×22 microlenses offering integral field spectroscopy over a field of 11.8×7.3 arcsec with a sampling of 0.52". The field covered by ARGUS was oriented north-east and was centered at the position $\alpha = 10:24:02.5$, $\delta = -57:45:24.0$ (see Fig. 1)

Table 1. Overview of the targets studied in this paper.

| Object | Inst. | Epochs | Binary status | Previous SpT | New SpT |
|-------------|---------------|--------|---------------|-------------------|-------------------|
| WR 20b | UVES | 8 | candidate | WN6ha | WN6ha |
| MSP 18 | UVES | 8 | candidate | O5-5.5 V-III((f)) | O5-5.5 V-III((f)) |
| MSP 44 | UVES | 8 | eclipsing | O9.5 V | B1 V + PMS? |
| MSP 96 | UVES | 8 | eclipsing | OB | B1: + B1: |
| MSP 167 | GIRAFFE-ARGUS | 3 | unknown | O6 III | earlier than O6 |
| MSP 171 | GIRAFFE-IFU | 5 | unknown | O4-5 | O4-5 |
| MSP 182 | GIRAFFE-IFU | 5 | unknown | O4 | O4 |
| MSP 183 | GIRAFFE-ARGUS | 3 | unknown | O3 V((f)) | O3 V((f)) |
| MSP 199 | GIRAFFE-ARGUS | 3 | unknown | O3-4 V | O3-4 V |
| MSP 203+444 | GIRAFFE-IFU | 5 | unknown | O6 V-III | O4-5 V |
| MSP 223 | UVES | 8 | eclipsing? | OB | O7-8 |
| A | GIRAFFE-ARGUS | 3 | unknown | — | O8 V |
| B | GIRAFFE-ARGUS | 3 | unknown | — | O6-7 |
| C | GIRAFFE-ARGUS | 3 | unknown | — | O7 V |
| D | GIRAFFE-ARGUS | 3 | unknown | — | O9.5: |
| E | GIRAFFE-ARGUS | 3 | unknown | — | O6-7: |

Notes. The first column lists the designation of the star, the second and third columns yield the instrument and the number of epochs for each target. The last three columns provide the binary status and spectral types prior to our study and the current spectral type (determined in this paper).

Table 2. Radial velocities of prominent interstellar absorption and nebular emission lines in the FLAMES-UVES spectra of our targets.

| Line | WR 20b (km s ⁻¹) | MSP 18 (km s ⁻¹) | MSP 44 (km s ⁻¹) | MSP 96 (km s ⁻¹) | MSP 223 (km s ⁻¹) |
|------------------------|---------------------------------|---------------------------------|---------------------------------|---------------------------------|----------------------------------|
| Na I λ 5890 | 3.4 \pm 0.3 | 2.8 \pm 0.3 | 2.8 \pm 0.5 | 2.2 \pm 0.2 | 2.3 \pm 0.4 |
| Na I λ 5896 | 3.7 \pm 0.2 | 2.0 \pm 0.5 | 2.9 \pm 0.5 | 1.8 \pm 0.3 | 3.2 \pm 0.2 |
| [O III] λ 4959 | 17.1 \pm 7.6 | -3.7 \pm 1.0 | 28.3 \pm 1.7 | 22.3 \pm 0.3 | 7.4 \pm 2.6 |
| [O III] λ 5007 | 18.6 \pm 1.7 | -3.1 \pm 1.5 | 29.7 \pm 1.2 | 22.6 \pm 0.8 | 8.2 \pm 0.9 |

Notes. The error bars correspond to the 1- σ dispersion in our measurements

near MSP 183 and MSP 167². For both runs (A and B), the GIRAFFE spectrograph was used in the low-resolution standard setup LR04 in third order covering the domain 5015 – 5815 Å with a measured resolving power of 9500.

All the data were reduced with the dedicated pipelines (versions 4.2.1 for UVES, 2.5.1 for GIRAFFE-IFU and 2.8.1 for GIRAFFE-ARGUS). The signal-to-noise ratio (S/N) of the individual FLAMES-UVES spectra reached between about 50 for the faintest stars and around 250 for WR 20b and MSP 18. For the three FLAMES-GIRAFFE spectra, the S/N was around 150 for MSP 171 and 182 and near 300 for MSP 203 + MSP 444. For the ARGUS data, the wavelength solution was obtained in two iterations yielding an rms in the calibration of 2.7 km s⁻¹. The spectra were extracted using the SUM option.

Owing to the high extinction towards Westerlund 2, the spectra of our targets are very rich in interstellar features. Besides the strong Na I $\lambda\lambda$ 5890, 5896 absorption lines, we observe a large number of diffuse interstellar bands (DIBs, see Herbig 1995). In addition, the spectra display a number of nebular emission lines from the RCW 49 complex, such as [O III] $\lambda\lambda$ 4959, 5007 as well as H I and He I.

As can be seen from the results quoted in Table 2, the interstellar absorption lines are consistently found to indicate radial velocities (RVs) of around 2 – 3 km s⁻¹ for each star, whilst the

radial velocities of the nebular emission lines span a range from -3.7 to +29.7 km s⁻¹, implying that the RCW 49 nebula has a rather complex velocity field.

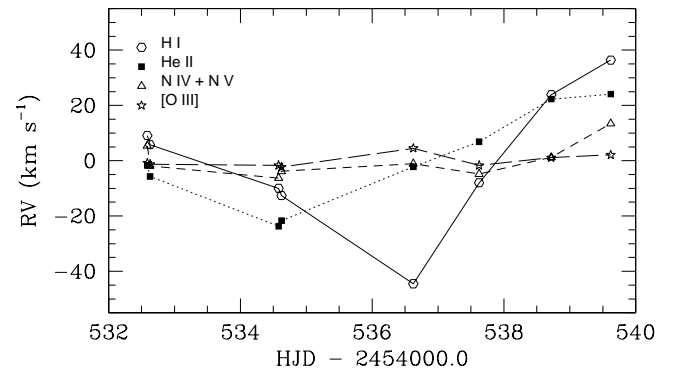


Fig. 2. Radial velocity variations about the mean evaluated from the five measurements for several sets of lines in the spectrum of WR 20b (see text). The RV measurements of the narrow nebular [O III] λ 5007 emission line are shown for comparison.

² The coordinates quoted in this paper are those from the archive WFPC2 HST image (based on GSC). They are shifted by -0.027 s in RA and -0.69" in Dec with respect to the 2MASS system used in Paper I.

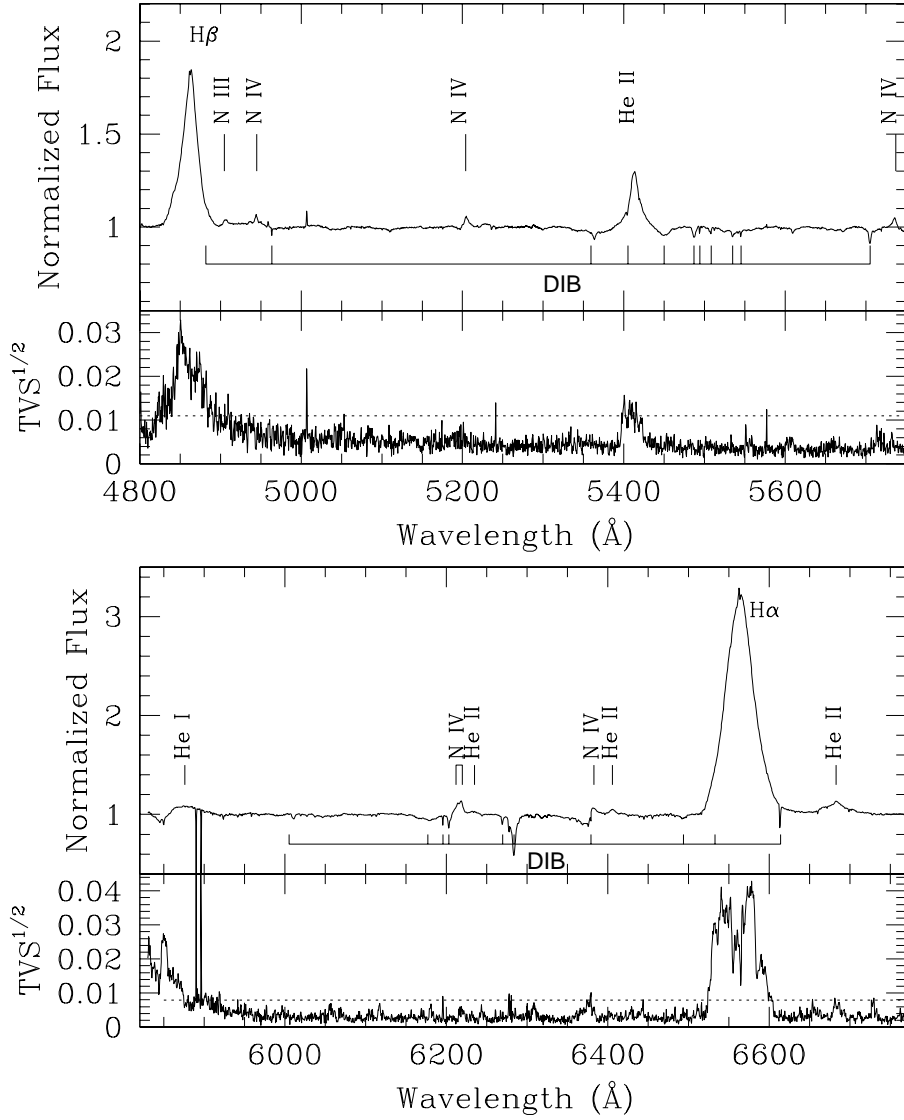


Fig. 3. Mean spectrum and temporal variance spectrum as computed from our time series of FLAMES-UVES spectra of WR 20b. The dashed lines yield the 99% significance level of the $\text{TVS}^{1/2}$.

3. Results

3.1. WR 20b

WR 20b was reported as an $\text{H}\alpha$ emission star by Thé (1966, his object THA 35-II-37). This result was confirmed by Schwartz, Persson & Hamann (1990), and the star was first classified as a WN7 Wolf-Rayet object by Shara et al. (1991). This classification was slightly revised to WN7:h, i.e. a WN7 star with a definite signature of hydrogen in its spectrum, by Shara et al. (1999), whilst van der Hucht (2001) proposed instead a spectral type WN6ha, i.e. a WN6 star with hydrogen emission lines and absorption lines belonging to the WN6 star. Subsequent radio observations revealed the presence of a wind-blown shell around WR 20b (Whiteoak & Uchida 1997). Whilst the mean photometric properties of WR 20b do not support the presence of a luminous companion, the X-ray spectrum of the star is remarkably hard, a feature that is often associated with interacting wind binaries (Nazé et al. 2008).

To address the issue of multiplicity, we measured the RVs of a number of stellar emission lines ($\text{H}\beta$, $\text{H}\alpha$, $\text{He II } \lambda\lambda 5412, 6406,$

6683 , $\text{N IV } \lambda\lambda 5204, 5737, 6383$, and $\text{N V } \lambda 4945$, see Table A.1), as well as interstellar and nebular features such as the diffuse interstellar bands (DIBs) at $\lambda\lambda 5363, 5705, 5850, 6010, 6614$ (see Herbig 1995), the $\text{Na I } \lambda\lambda 5890, 5896$ doublet, and the $[\text{O III}] \lambda 5007$ line. The interstellar and nebular features allow us to check the stability of the wavelength calibration: their $1 - \sigma$ RV dispersion is 0.3 km s^{-1} for the narrow Na I absorptions, indicating a very stable calibration. For the broad DIBs, the dispersion is larger, of the order of 4 km s^{-1} . For the majority of the stellar emission lines, the RV dispersions are significantly larger than these values ($24.8, 17.7$, and 6.5 km s^{-1} for the H I , He II and $\text{N IV} + \text{v}$ lines, respectively). However, as shown in Fig. 2, the variations in the different sets of lines are uncorrelated.

To quantify the spectral variability of the star, we computed the temporal variance spectrum (TVS, see Fullerton, Gies & Bolton 1996) of our time series. This method provides a quantitative assessment of the significance of the spectral line profile variability properly accounting for the noise level of the individual spectra. The results are shown in Fig. 3. Significant variations are detected in the strongest emission lines, but not in

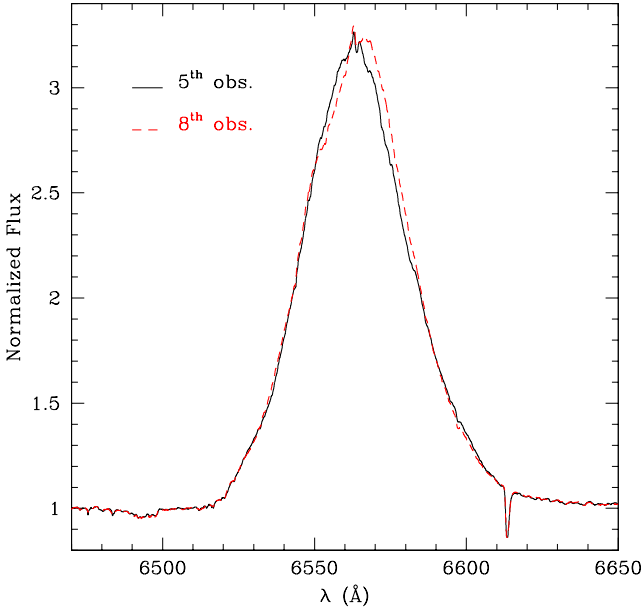


Fig. 4. Illustration of the $H\alpha$ line profile in the spectrum of WR 20b for those two observations corresponding to the maximum difference in apparent RV. The RV variations are clearly due to the top part of the line profile only.

weaker ones such as the $N\text{IV } \lambda\lambda 6212\text{--}20$ blend. The TVS of the strongest lines is quite complex, with multiple (up to four) peaks, indicating that the RV variability of these lines reflects their line profile variability. A more detailed inspection of the profiles of the hydrogen lines in different observations (see Fig. 4) actually indicates that the apparent RV variations are mainly due to changes in the upper part of the line profiles. The results presented in this section suggest that WR 20b is indeed a spectroscopically variable object, but this variability is more likely to be related to (large-scale) structures in the stellar wind, than caused by genuine RV variations of the star. Therefore, we conclude that there is at present no firm evidence that WR 20b is a short period (less than two week period) spectroscopic binary system (see also the discussion in Sect. 4).

Following the classification criteria of Smith, Shara & Moffat (1996), our spectra correspond most closely to a WN5-6 spectral type (ratio of the $\text{He II } \lambda 5412$ to $\text{He I } \lambda 5876$ intensities above continuum = 3.6) with narrow lines (equivalent width, EW, of $\text{He II } \lambda 5412 = -5.9 \text{ \AA} > -40 \text{ \AA}$) and a definite signature of hydrogen emission revealed by the oscillation of the Pickering decrement (compare the intensities of $H\alpha$, $\text{He II } \lambda 5412$, and $H\beta$ on Fig 3). The only obvious stellar absorption feature in our spectra is the P-Cygni-type absorption trough of the $\text{He I } \lambda 5876$ line, which is obviously associated with the Wolf-Rayet star. Therefore, our data are consistent with the WN6ha classification given by van der Hucht (2001).

3.2. MSP 18

On the basis of low-resolution spectroscopy, Moffat et al. (1991) classified MSP 18 as an O7:V star. This classification was revised to O4 V((f)) by Uzpen et al. (2005) and more recently to O5-5.5 V-III((f)) based on higher resolution spectra (Paper I). Both its optical brightness (Paper I) and X-ray emission properties (brightness, spectral hardness and slight flux variability; see

Table 3. Radial velocities of MSP 96.

| Date HJD-2454000 | ϕ_p | ϕ_{RV} | Primary (km s^{-1}) | Secondary (km s^{-1}) |
|---------------------|----------|-------------|-----------------------------------|-------------------------------------|
| 532.587 | 0.044 | 0.058 | | 154 |
| 532.631 | 0.085 | 0.099 | -182 | 159 |
| 534.581 | 0.903 | 0.917 | 239 | -113 |
| 534.628 | 0.946 | 0.960 | 112 | -80 |
| 536.624 | 0.807 | 0.821 | 289 | -230 |
| 537.626 | 0.741 | 0.755 | 374 | -259 |
| 538.719 | 0.760 | 0.774 | 352 | -219 |
| 539.622 | 0.602 | 0.616 | 195 | -167 |

Notes. The RVs are inferred from the multi-Gaussian deblending of the $\text{He I } \lambda 6678$ line accounting for a nebular emission line. The second and third columns yield the orbital phases computed according to the photometric and spectroscopic ephemerides, respectively.

Nazé et al. 2008) suggest that this star could be an interacting-wind binary system.

We fitted Gaussian line profiles to measure the RVs of a number of stellar absorption lines ($H\beta$, $H\alpha$, $\text{He I } \lambda 5876$, $\text{He II } \lambda\lambda 5412, 6683$), as well as two faint emission lines ($\text{C III } \lambda\lambda 5696, 6729\text{--}6731$; Walborn 2001). The results are listed in Table A.2. We also measured the same interstellar and nebular features as for WR 20b as well as the nebular $H\beta$ and $H\alpha$ emission lines. The $1 - \sigma$ RV dispersions of the stellar absorption lines are between 1.8 and 5.9 km s^{-1} , which is quite comparable to the corresponding values of the prominent interstellar features. The stellar emission lines have slightly larger RV dispersions, but these emission lines are actually extremely weak (normalized intensity of 0.02 above the continuum for $\text{C III } \lambda 5696$) and their RVs are thus subject to larger uncertainties. In addition, the TVS does not reveal any significant line profile variability. We thus conclude that MSP 18 is unlikely to be a short period (shorter than two week) spectroscopic binary (see also the discussion in Sect. 4).

3.3. MSP 96

The star MSP 96 was found to be an OB-type eclipsing binary system with a period of (1.0728 ± 0.0006) days with probably at least one component close to filling its Roche lobe (Paper I). Our spectra do not reveal any trace of the $\text{He II } \lambda 5412$ line, thus they are most consistent with a spectral type later than B0.2 (Walborn 1980). As a result of the severe contamination of the spectra by strong and asymmetric nebular emission lines in the hydrogen and neutral helium transitions, measuring the stellar components in the spectra of MSP 96 is quite a challenging task.

The stellar spectral lines are broad, as expected for rapidly rotating stars in a contact binary system. Using a multi-Gaussian fitting technique, we attempted to deblend the $H\alpha$, $\text{He I } \lambda\lambda 5871, 6678$ lines, accounting for the presence of a strong nebular emission line. Our most reliable results were obtained for the $\text{He I } \lambda 6678$ line and the corresponding RVs are listed in Table 3. Unfortunately, the uncertainties on these RVs are quite large, of the order of 20 km s^{-1} . The He I lines are essentially identical in strength for the spectra of both stars, whilst the $H\alpha$ line appears to be up to four times stronger in the spectrum of the star that has negative velocities during most of our observations. We caution however that this latter result is probably an overestimate because of a bias resulting from the presence of the strong nebular emission line.

In our data, the two components are most clearly deblended for the fifth, sixth, and seventh observations. Extrapolating the

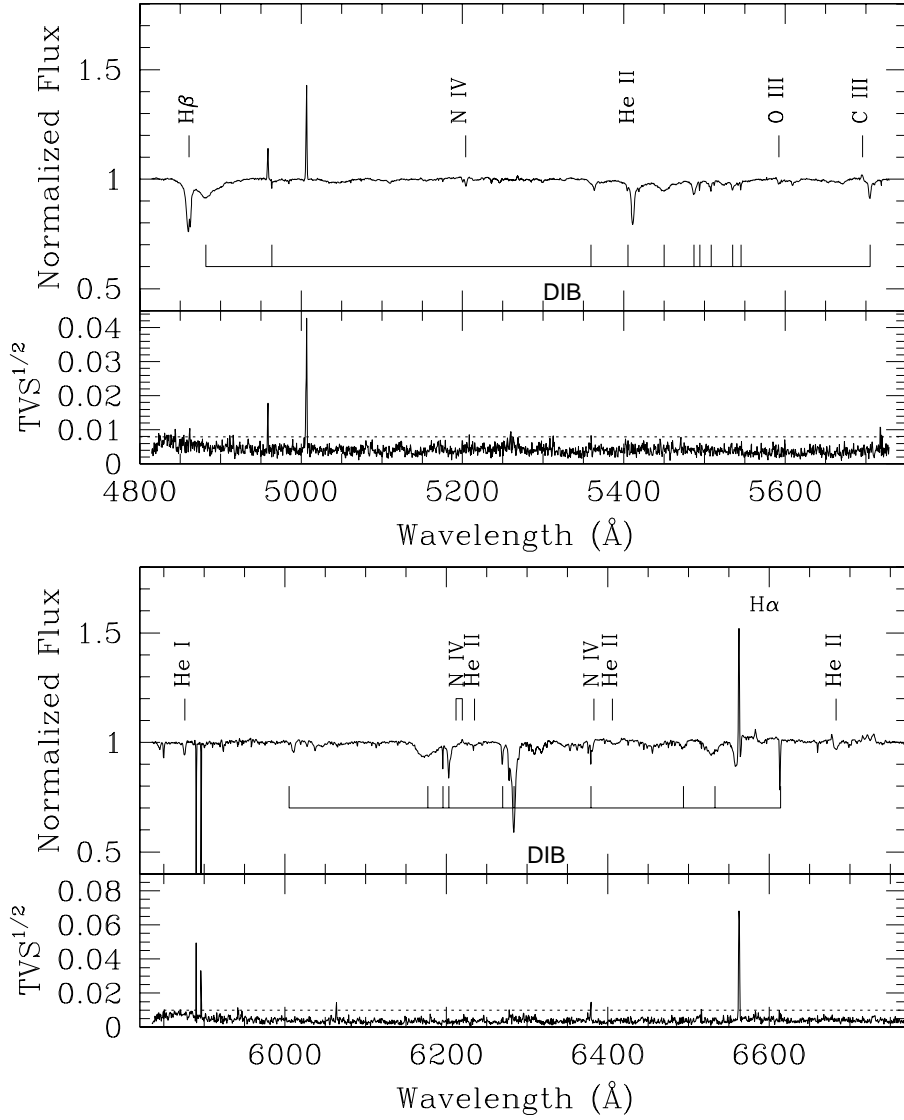


Fig. 5. Same as Fig. 3, but for MSP 18.

light curve ephemerides of Paper I to the epoch of our spectroscopic observations, we find that these observations should correspond to orbital phases near $\phi_p = 0.75$ (where $\phi_p = 0.0$ corresponds to the primary eclipse, i.e. the brighter star being eclipsed). Although our spectroscopic data are taken 1162 days after T_0 of Paper I and despite the short orbital period, the orbital solution below agrees with the photometric ephemerides. Indeed, the shift in phase between the photometric and spectroscopic solution is found to be 0.014.

We note that the star with the stronger $H\alpha$ line is seen blueshifted at phases near $\phi_p = 0.75$, hence that this line belongs to the intrinsically fainter component (i.e. the photometric secondary) of the system. From these considerations, we conclude that the photometric secondary must be intrinsically somewhat cooler than the primary, but should have a somewhat larger radius. The helium line intensity ratio might then be close to unity, although the primary star (i.e. the one that is eclipsed at phase 0.0) would be somewhat hotter, thus explaining the deeper primary eclipse. Alternatively, the star with the negative RVs near $\phi_p = 0.75$ could be the primary star, but this would imply a shift in phase between the photometric and spectroscopic

ephemerides of 0.486. This seems *a priori* rather unlikely, but cannot entirely be ruled out given the error in the orbital period and the large number of orbital cycles between our photometric and spectroscopic campaigns. We briefly consider the implications of this alternative assumption below.

3.3.1. Orbital solution

To establish the orbital solution of MSP 96 from the RV data in Table 3, we started by considering the SB1 solution for the photometric secondary star. For this purpose, we used the Liège Orbital Solution Package (LOSP) code (Sana & Gosset 2011), which allowed us to derive $K_s = (270 \pm 22) \text{ km s}^{-1}$, as well as $\gamma = (19 \pm 19) \text{ km s}^{-1}$. We then considered the RVs of the photometric primary star with the value of γ fixed to 19 km s^{-1} , which yielded $K_p = (337 \pm 78) \text{ km s}^{-1}$. From these results, we then inferred a secondary/primary mass ratio of 1.25 ± 0.27 , as well as $m_p \sin^3 i = (11.1 \pm 3.0) M_\odot$ and $m_s \sin^3 i = (13.8 \pm 6.6) M_\odot$. These minimum masses, along with the eclipsing nature of the system, which indicates a rather large inclination, suggest a spectral type near B1 for both components. This classification is also

consistent with the overall spectral morphology, although the large number of interstellar features and the moderate S/N of our data prevent us from being more specific.

3.3.2. Light curve solution

Adopting the mass ratio inferred hereabove, we then analysed the photometric data of the system from Paper I using the NIGHTFALL³ code, developed and maintained by R. Wichmann, M. Kuster and P. Risse. We set the temperature of the primary star to 25400 K (appropriate for a B1 V classification, Schmidt-Kaler 1982) and allowed the inclination, the Roche-lobe filling factors (hereafter $fill_p$ and $fill_s$), and the secondary temperature to vary. The best-fit model parameters were found to be $i = (76.8^{+0.1}_{-0.2})^\circ$, $T_s = 21550^{+120}_{-260}$ K, $fill_p = 0.84 \pm 0.03$, and $fill_s = 0.85 \pm 0.02$. This inclination yields absolute masses and radii, respectively of $M_p = (12.0 \pm 3.3) M_\odot$, $M_s = (15.0 \pm 7.2) M_\odot$, and $R_p = (3.8 \pm 0.6) R_\odot$, $R_s = (4.3 \pm 0.6) R_\odot$. These radii are rather small for the masses, but we note that quite similar parameters were found for other early B-type main-sequence stars in eclipsing binaries (e.g. V 578 Mon and DW Car, see Southworth 2007 and references therein). That the more massive star (the secondary) would be the cooler and less luminous component of the system is surprising and indicates that this star would actually be more evolved. The radii of both stars suggest instead that they are rather unevolved main-sequence stars. Whilst a case A binary mass transfer could alter the properties of the components, we stress that the light curve solution relates to stars that do not fill up their Roche lobes. This issue would actually be resolved if we assumed a phase shift of 0.486 between the photometric and spectroscopic ephemerides (see below).

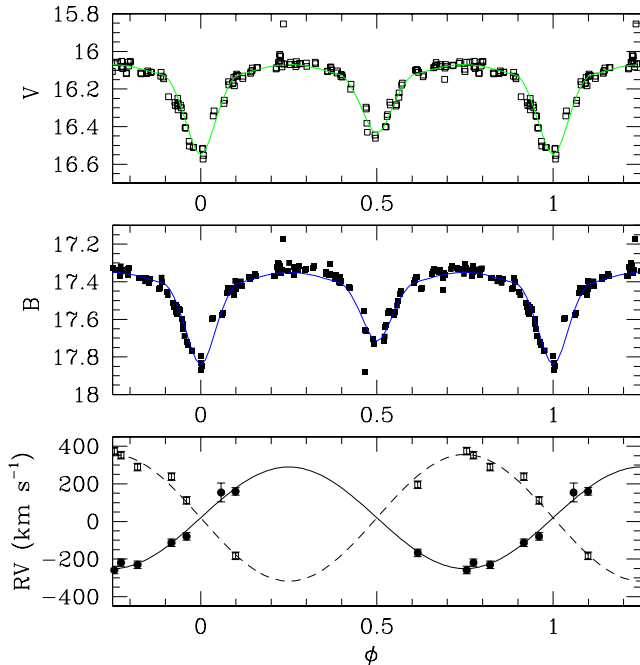


Fig. 6. Combined solution of the light curve and the RV curve of the MSP 96 system (see text). The orbital phases are computed from the photometric ephemerides of Paper I.

Combining all the above parameters, we inferred a total bolometric luminosity of the binary of $8960 L_\odot$. The uncertainty in this number is difficult to evaluate, because of the unknown uncertainty in the primary's spectral type. However, we estimated that the relative uncertainty is likely to be about 25%. From $V = 16.17$ at phases outside the eclipses and $E(B - V) = 1.50$ as follows from Fig. 7 of Paper I, we then inferred a distance of 6.6 kpc again probably subject to a 25% uncertainty.

We also analysed the light curve assuming a shift of 0.486 in phase between the photometric and spectroscopic ephemerides. In this case, the mass ratio inferred in Sect. 3.3.1 would be inverted, $m_s/m_p = 0.80 \pm 0.17$, and the more massive star (which is now the primary with $m_p \sin i = 13.8 M_\odot$) would also be the brighter one. The corresponding best-fit photometric solution is essentially symmetric with respect to the parameters indicated above ($i = 77.2^\circ$, $T_p = 25400$ K (fixed), $T_s = 21790$ K, $fill_p = 0.85$, $fill_s = 0.84$) and yields a similar distance of 6.8 kpc. This solution corresponds to an unevolved B + B main-sequence binary system and avoids the problems described above.

3.4. MSP 44

In Paper I, we classified MSP 44 as an eclipsing binary with a period of 5.176 days, noting that the aliasing period at 1.2395 days yielded an almost equally strong peak in the periodogram. However, the rather narrow eclipses argue against the shorter orbital period and we thus favour the solution corresponding to the (5.176 ± 0.029) day period. Extrapolating the photometric ephemerides from Paper I to the epoch of our FLAMES observations, we find the phases listed in the second column of Table 4, where phase 0.0 corresponds to the primary eclipse.

Whilst the primary eclipse has a depth of 0.6 mag, the secondary eclipse was found to be much shallower (< 0.1 mag). There are two possible interpretations of this situation: either the surface brightnesses of the two components of the system are very different or the binary orbit is sufficiently eccentric with a longitude of periastron (for the star being eclipsed at phase 0.0) near 90° ⁴. In addition, the light curve exhibits some slow modulation, especially around secondary minimum that we tentatively attributed to spots on one of the stars (Paper I).

The absence of any He II line in the spectrum of MSP 44 clearly indicates a spectral type later than B0.2 (Walborn 1980) for the combined spectrum of MSP 44. This is slightly at odds with the O9.5V: photometric spectral classification of this star (Nazé et al. 2008). The likely presence of a very weak Si III λ 5740 feature and possibly of Si II λ 6347 suggest a spectral type near B1 with an uncertainty of at least one subtype. Given the probably rather large brightness ratio of the primary to the secondary, we can actually adopt this spectral type as the one of the primary component. The B1 spectral classification is also surprising in view of the rather strong and hard X-ray emission of MSP 44 measured with *Chandra* (Nazé et al. 2008).

Measuring the stellar RVs was, once more, quite difficult. The reasons for this were (1) the contamination of most, sufficiently intense, stellar lines (H I and He I) by rather strong nebular emission lines, and (2) the moderate S/N of our data that

⁴ If the eccentricity were sufficiently large and the longitude of periastron close to 90° , one could expect to observe a difference in the durations of the primary and secondary eclipses, although the amplitude of this effect strongly depends on a number of parameters (orbital eccentricity and inclination, filling factors,...) and is sometimes even absent. Within the accuracy of our data, no such effect is present in the light curve of MSP 44.

³ <http://www.hs.uni-hamburg.de/DE/Ins/per/Wichmann/Nightfall.html>

Table 4. Radial velocities of MSP 44.

| HJD -2454000 | ϕ_p | ϕ_{RV} | RV ₁ (km s ⁻¹) | He I λ 5876 RV _b (km s ⁻¹) | RV _r (km s ⁻¹) | He I λ 4921 RV ₁ (km s ⁻¹) |
|-----------------|----------|-------------|--|---|--|---|
| 532.587 | 0.273 | 0.157 | 17.3 | -285 | | 5.3 |
| 532.631 | 0.281 | 0.166 | 26.4 | -269 | | 25.6 |
| 534.581 | 0.658 | 0.542 | 45.9 | -252 | 309 | 59.8 |
| 534.628 | 0.667 | 0.551 | 36.6 | -267 | 287 | 66.5 |
| 536.624 | 0.053 | 0.937 | | -246 | | 46.5 |
| 537.626 | 0.247 | 0.131 | 27.0 | -273 | | 12.1 |
| 538.719 | 0.458 | 0.342 | | | | 29.6 |
| 539.622 | 0.632 | 0.516 | 45.1 | -264 | 315 | 27.6 |

Notes. The second and third columns yield the orbital phases computed according to the photometric and spectroscopic ephemerides respectively. RV₁, RV_b, and RV_r stand for the RVs of the main component of the He I λ 5876 line and the blue and red components of this line, respectively (see text).

hampers the detection of very faint spectral lines. At first sight, only the signature of one star is seen in the spectra. The RVs of this main component vary very little during our observations (which cover *a priori* both quadrature phases of the orbit). This is very puzzling and prompted us to re-inspect the morphology of several lines in more detail.

The most intriguing line turned out to be He I λ 5876. The blue and red wings of this line suggest the presence of two additional absorption features in most of our data (see Fig. 7). There are no known DIBs at these wavelengths (Herbig 1995)⁵ and the blue wing of the line at least should not be affected by any strong telluric absorption lines either (Curcio, Drummer & Kneistrick 1964).

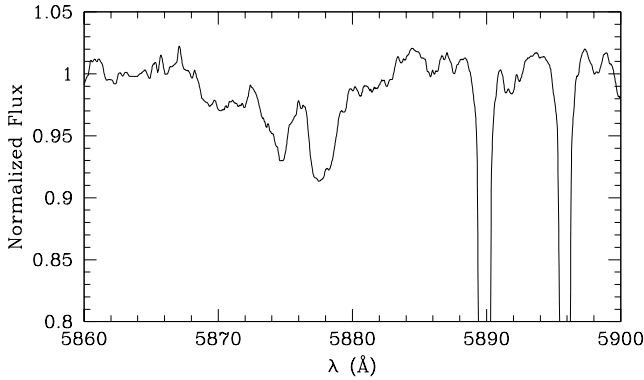


Fig. 7. The region around the He I λ 5876 line in the spectrum of MSP 44 observed on HJD 2454539.622. The emission peak near 5876 Å is the nebular He I line. Note the complex morphology of the line, especially in the blue and red wing.

Since it is possible that MSP 44 might be a triple system, we thus performed a multiple Gaussian fit of this line with up to four components (three stellar absorption components and one nebular emission component). The nebular component was consistently fitted with an RV of (29.4 ± 3.1) km s⁻¹. The main stellar component was in turn found to display RVs between 17 and 46 km s⁻¹. The corresponding RVs are listed in Table 4. Although the RVs of the blue- and red-shifted absorption features must be interpreted with extreme caution, it is immediately

clear that if MSP 44 were a triple system, then the mass ratio of the binary responsible for the blue and red absorption components would be close to unity. This in turn would be at odds with the light curve, which implies that there is a rather large difference in surface brightness between the components of the eclipsing binary. The most likely interpretation of this line morphology therefore seems that it represents the very broad absorption line of only one star, but either deformed by pulsations such as in β Cep-type variables (e.g. Telting et al. 2006) or partially filled in by some circumstellar emission or chromospheric emission from a low-mass companion heated by the irradiation from the B-star. In the case of the former interpretation, it must be stressed that our photometry does not reveal the signature of typical β Cep variations (see Paper I). As to the circumstellar emission interpretation, we note that neither H β nor H α have emission above the continuum, but both lines have a somewhat similar morphology to that of He I λ 5876, in the sense that they exhibit a rather narrow core and more extended wings (especially the red wing).

In the absence of an unambiguous interpretation of the spectral morphology of MSP 44, we choose to use the RVs of the He I λ 4921 line, which is less complex than He I λ 5876, to derive some constraints on the orbital motion of the main component. Assuming that the RVs of this line indeed reflect orbital motion, a circular SB1 orbital solution yields $T_0 = 2454531.774 \pm 0.285$, $\gamma_1 = (42 \pm 9)$ km s⁻¹, $K_1 = (34 \pm 16)$ km s⁻¹, and $a_1 \sin i = 3.5 R_\odot$. This yields a very low mass-function of $f(m) = \frac{m_2^3 \sin^3 i}{(m_1 + m_2)^2} = (0.021 \pm 0.030) M_\odot$. Here, T_0 stands for the time of primary minimum. Compared to the photometric ephemerides, there is a shift in phase by 0.116 or 1.116. Since the number of orbital cycles between our photometric and spectroscopic campaigns amounts to 224, this phase shift implies an error in the orbital period of either 0.003 or 0.026 days, well within or consistent with the estimated error in the orbital period as inferred from the photometric data (0.029 days, see above).

Since the system is eclipsing, we can to first order approximate the $\sin^3 i$ term in the mass function to be unity. Assuming a mass of $10 - 12 M_\odot$ for the B1 primary, we find that the secondary should have a mass of about $1.4 - 1.6 M_\odot$, which implies that the companion might be an early F-type star. But would this result be reasonable? To answer this question, we return to the photometric light curve of the system. We first recall that the latter displays variations outside the eclipses that are quite large. In Paper I, we attributed them to spots, but our current knowledge of the system allows us to be a bit more specific. If we are indeed dealing with an F star closely orbiting around a B1 V primary, it

⁵ An up-to-date catalog of DIBs can be found at <http://leonid.arc.nasa.gov/DIBcatalog.html>.

seems unavoidable that there will be a strong heating effect of part of the secondary's surface by the primary's radiation. Test calculations carried out with the NIGHTFALL code indicate that the light curve of the system can *a priori* be explained by assuming that the hemisphere of the secondary facing the primary is significantly hotter than the rear side. However, the solution of the light curve is certainly not unique, because we ignore the actual mass ratio of the system and this parameter cannot be well-constrained from the photometric data alone. As an illustration, we show below the best-fit solution obtained for $m_1/m_2 = 8$.

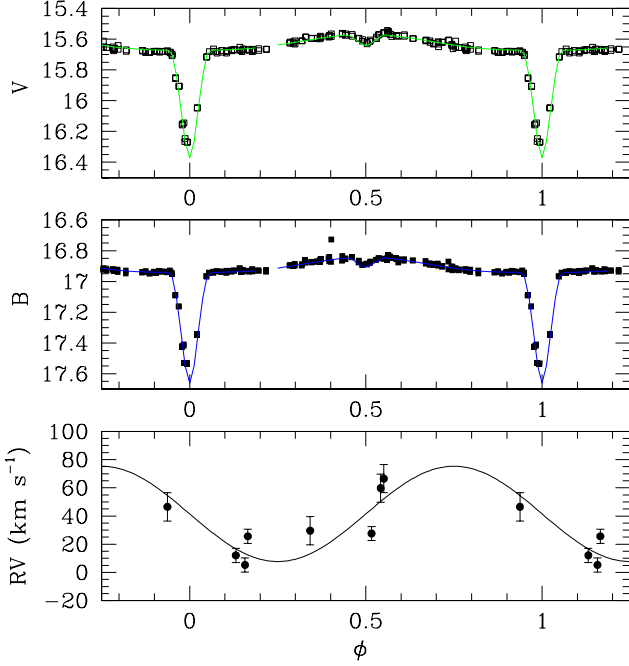


Fig. 8. Combined solution of the light curve and the SB1 RV curve of MSP 44. The parameters of the light curve model are $m_2/m_1 = 0.125$, $i = 80.6^\circ$, primary and secondary radii of 5.2 and $5.6 R_\odot$ respectively, and mean surface temperatures of $T_1 = 29100$ K (fixed) and $T_2 = 6440$ K. The secondary star features a hot spot facing the primary with a radius of 45° and a temperature enhanced by a factor 1.76.

We note in addition that the radius of the secondary inferred from the light curve solution ($5.6 R_\odot$ in the case of the best-fit solution shown below) is actually much larger than the expected radius of an early-F main-sequence star, but closer to what one would expect for a giant. Given the youth of the Westerlund 2 cluster, the most straightforward explanation is that the secondary star of MSP 44 is not a genuine giant, but rather a pre-main sequence (PMS) object. A PMS nature of the secondary would also provide a very natural explanation of the rather bright and hard X-ray emission of this system.

In summary, our analysis of the spectroscopic and photometric data of MSP 44 leaves a number of open questions about the exact nature of the primary component. The most likely scenario seems to be that this system consists of a B1 main-sequence star and a low-mass PMS companion. The physical parameters of the system are admittedly too uncertain to attempt a self-consistent determination of the distance of Westerlund 2, although we note that the parameters of the fit in Fig. 8 are consistent with a distance of 8.4 kpc. If the B1 V spectral type of the MSP 44 primary were correct and the companion contributed little to the

integrated light of the system, then the V magnitude and $B - V$ colour outside eclipse would indeed be consistent with a distance near 8.4 kpc.

3.5. MSP 223

The yellow-red spectra of MSP 223 display a rather strong He II $\lambda 5412$ feature, as well as a weak ($EW \sim 0.4 \text{ \AA}$), but definitely detected O III $\lambda 5592$ line (see Fig. 9). Whilst there are no known interstellar features affecting the O III line, there are three, rather weak, DIBs at $\lambda\lambda 5404.5$, 5414.8 , and 5420.2 that are blended with He II $\lambda 5412$. To assess the real strength of the He II $\lambda 5412$ feature, we first evaluated the ratio of the EWs for some nearby DIBs ($\lambda\lambda 5450.3$, 5487.5) in the spectrum of MSP 223 and the corresponding EWs in the spectrum of HD 183143 quoted by Herbig (1995). This ratio (1.33) was then used as a linear scaling factor to estimate the contamination of the He II $\lambda 5412$ line by the three DIBs. The corrected EW of the He II line amounts to 0.9 \AA . Therefore, He II $\lambda 5412$ is of comparable strength to the He I $\lambda 5876$ line ($EW = 0.85 \text{ \AA}$), although the latter is slightly affected (probably less than 0.2 \AA) by a nebular emission line. These properties suggest an O7-8 spectral type (Walborn 1980) for the combined spectrum. That we are dealing with an O-type star is confirmed by the absence of He I $\lambda 4921$ and the fact that the He I $\lambda 6678$ line is clearly blended with He II $\lambda 6683$.

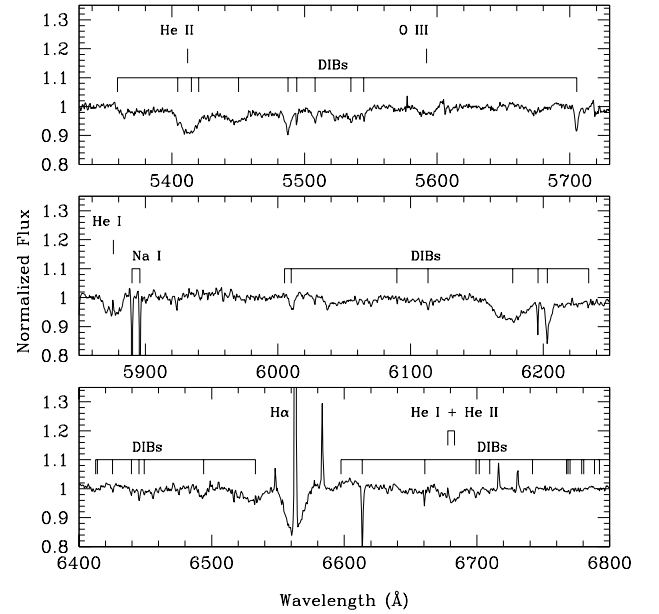


Fig. 9. Mean spectrum of MSP 223. Most of the features are due to interstellar lines or bands. The rather prominent He II $\lambda 5412$ line indicates an O-star classification (see text for details).

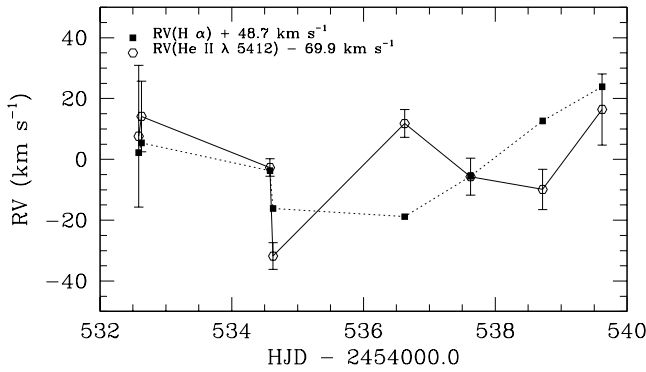
We measured the RVs of several stellar lines. The most robust results were obtained for He II $\lambda 5412$ and H α , although we caution that the former is affected by blends with DIBs and for the latter, we had to perform a two Gaussian fit to account for the contamination by the nebular H α emission line⁶. The results are listed in Table 5 and plotted in Fig. 10, after subtracting the mean RV individually for each line. No obvious trend is apparent in this figure. The rather large shift in RV between the He II

⁶ In these fits, the nebular line has a very stable RV of $1.1 \pm 0.2 \text{ km s}^{-1}$.

Table 5. Radial velocities of MSP 223.

| HJD-2454000 | RV(He II λ 5412) (km s ⁻¹) | RV(H α) (km s ⁻¹) |
|-------------|---|--|
| 532.587 | 77.5 | -46.4 |
| 532.631 | 84.0 | -43.2 |
| 534.581 | 67.2 | -52.4 |
| 534.628 | 38.1 | -64.8 |
| 536.624 | 81.7 | -67.5 |
| 537.626 | 64.2 | -54.2 |
| 538.719 | 60.0 | -36.0 |
| 539.622 | 86.3 | -24.8 |

λ 5412 and H α line is most likely due to wind emission affecting the latter line. Similar shifts in the apparent systemic velocity of different lines are commonly found in binary systems where at least one component features a strong stellar wind (e.g. Rauw et al. 2002).


Fig. 10. Radial velocity variations of MSP 223 during our observations. The error bars in the H α RVs yield the formal error in the position of the absorption line in the two-Gaussian fit accounting for the nebular emission line.

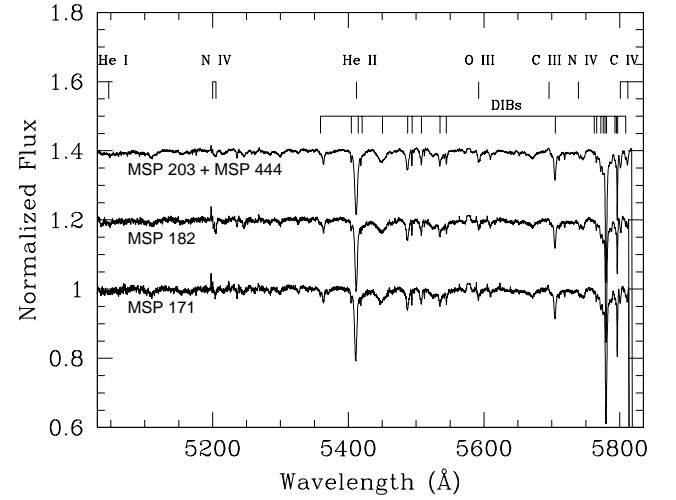
Our previous photometric monitoring of the cluster indicated that MSP 223 is probably an eclipsing binary (Paper I). The Fourier analysis yielded the highest peak for a period of 14.61 days, with a rather large uncertainty of 1.2 days. The light curve folded with this period yields only a primary eclipse, with no indication of a secondary eclipse, although we might have missed it, if its duration is sufficiently short to fit into a small gap in phase coverage around $\phi = 0.5$. In Paper I, we inferred an orbital period of 29.211 days, which then indicates two rather similar eclipses. This would imply that the system consists of two rather similar stars, well within their Roche lobe and seen at a large inclination angle. However, we do not observe an obvious SB2 signature in our spectra of MSP 223, which suggests instead that the companion of the O-star is significantly fainter and less massive. This would be consistent with the rather modest RV variations that we have measured (see Fig. 10), which show a peak-to-peak amplitude of about 40 km s⁻¹ over seven days (the most reliable RVs are probably those of H α). Unfortunately, the considerable uncertainties in the orbital period prevent us from extrapolating the photometric ephemerides to the epoch of our spectroscopic observations and we cannot derive constraints on the orbital solution of this system.

Although our knowledge about MSP 223 is still rather poor, we can nevertheless attempt to derive its spectro-photometric distance. For this purpose, we combine the photometric informa-

tion from Paper I ($V = 15.75$, $B - V = 1.50$) with the absolute magnitudes and intrinsic colours of O7-8 V stars from Martins & Plez (2006). In this way, we obtain a distance modulus of 14.67 – 14.96 corresponding to a distance of 8.6 – 9.8 kpc. If our spectral classification were too early, and we adopted instead an O9.5 V type, we would still derive a distance of 7.2 kpc. We stress that this number does not account for the contribution of a companion to the light of MSP 223. Accounting for this would lead us to inferring an even larger distance.

3.6. MSP 171, 182, and 203 + 444

These three (actually four) objects were observed with the integral field spectroscopy units connected to the GIRAFFE spectrograph. Their mean spectra are shown in Fig. 11.


Fig. 11. Mean GIRAFFE spectra of MSP 171, 182, and 203 + 444.

In Paper I, the combined blue spectrum of MSP 203 + MSP 444 was assigned a spectral type O6 V-III. However, the GIRAFFE spectra of MSP 203 + MSP 444 did not reveal any trace of He I $\lambda\lambda$ 5016, 5048. Therefore, by comparison with the yellow-green spectra of OB stars published by Kerton, Ballantyne & Martin (1999), the brightest star of the blend must have a spectral type earlier than O6 V and possibly even earlier than O5 V. The mean spectrum of MSP 203 + MSP 444 is very similar to that of MSP 171 and 182, which were classified as O4-5 and O4 respectively in Paper I. We thus conclude that the brighter component of MSP 203 + MSP 444⁷ is likely of spectral type O4-5 V. The other component is likely responsible for the weak He I λ 4471 absorption that was seen in Paper I. We note that the latter absorption is significantly broader than the other stellar absorption lines in the blue (and yellow-green) spectra of MSP 203 + MSP 444. This supports our suggestion that this line is not associated with the same star as the stronger, higher ionization lines.

We measured the RVs of the three stars by fitting Gaussian profiles to the He II λ 5412, O III λ 5592, and C IV $\lambda\lambda$ 5801, 5812 lines with the caveat that the C IV lines (especially C IV λ 5812) are heavily blended with DIBs. We also measured the RV disper-

⁷ According to the photometry of Moffat et al. (1991), MSP 203 is slightly brighter than MSP 444. However, the HST data (see Fig. 1) suggest the reverse situation.

sions of four DIBs ($\lambda\lambda$ 5362, 5487, 5705, and 5781, see Herbig 1995), which are of comparable strength to or slightly stronger than the stellar lines.

For MSP 171, we infer a $1 - \sigma$ RV dispersion of 6.9 km s^{-1} for the strongest stellar line (He II λ 5412) and 3.5 km s^{-1} for the C IV λ 5812 line. For the same set of spectra, the RV dispersions of the DIBs are between 0.8 and 10.3 km s^{-1} with a typical value of about 5 km s^{-1} .

In the case of MSP 182, all stellar lines, except for C IV λ 5812, which is affected by the DIB at 5809 \AA (Herbig 1995), display a $1 - \sigma$ RV dispersion in the range $3.0 - 5.7 \text{ km s}^{-1}$. For the interstellar bands, these dispersions are $1.1 - 4.1 \text{ km s}^{-1}$, except for the rather weak DIB at λ 5362 which has a larger dispersion of 9.3 km s^{-1} .

Finally, the strongest stellar lines of the MSP 203 + MSP 444 spectra have RV dispersions of $2.5 - 4.1 \text{ km s}^{-1}$, whilst the weak O III λ 5592 feature has a slightly larger dispersion of 7.6 km s^{-1} . In this case, the RV dispersions of all DIBs measured amount to $1.6 - 3.3 \text{ km s}^{-1}$.

In conclusion, none of the stars observed with the IFUs feeding GIRAFFE is found to display significant RV variations on timescales of a week.

3.7. Stars observed with the ARGUS unit

The field of view covered by the ARGUS unit harbours two stars with an MSP identifier (MSP 167 and 183) as well as five anonymous objects (labeled A to E in Fig. 1) that are bright enough to extract at least an averaged spectrum. In addition, one spaxel at the edge of the field contains photons from the wing of the PSF of star MSP 199 which is right outside the nominal field of view. For the brightest and most isolated objects, the spectra were extracted over nine spaxels (MSP 167, MSP 183, object A), whilst the extraction region was reduced to a single spaxel for those stars (objects B, C, D, and E) that are either relatively close to a bright source or fall only partially inside the field of view of the ARGUS unit (MSP 199). In this section we discuss the results obtained for the different stars in the ARGUS field.

In Paper I, MSP 183 was classified as an O3 V(f) star. The lack of He I $\lambda\lambda$ 5015, 5047 absorption, the strength of the He II λ 5412 line and the presence of the N IV $\lambda\lambda$ 5200 - 5205 blend in the GIRAFFE-ARGUS spectrum of MSP 183 (see Fig. 12) are in line with this classification (Kerton et al. 1999). However, the most remarkable feature of the spectrum of MSP 183 is the presence of the C IV $\lambda\lambda$ 5801, 5812 lines in weak, but definite, emission with equivalent widths of -0.24 \AA and -0.06 \AA , respectively. These lines were reported to be in emission in the spectrum of the O3 If* supergiant Cyg OB2 #7 (Walborn & Howarth 2000). Our blue spectrum of MSP 183 (see Paper I) rules out the possibility that MSP 183 could be such an evolved supergiant. Therefore, it seems more likely that these C IV emission lines are not selective properties of O3 If* supergiants, but could actually be a general feature of the hottest O-type stars as suggested by Walborn (2001).

The RVs of MSP 183 were measured by fitting Gaussian profiles to the strongest stellar lines (He II λ 5412, O III λ 5592, C IV $\lambda\lambda$ 5801, 5812; see Table A.6). In this way, we obtain a $1 - \sigma$ RV dispersion of 1.2 km s^{-1} for the He II λ 5412 absorption and 7.3 and 3.3 km s^{-1} for the C IV λ 5801 and λ 5812 emission lines respectively. The RV dispersion of the DIBs range between 0.9 and 7.5 km s^{-1} depending on the strength and shape of the DIB. We therefore conclude that this star does not display any significant RV variations during our observations.

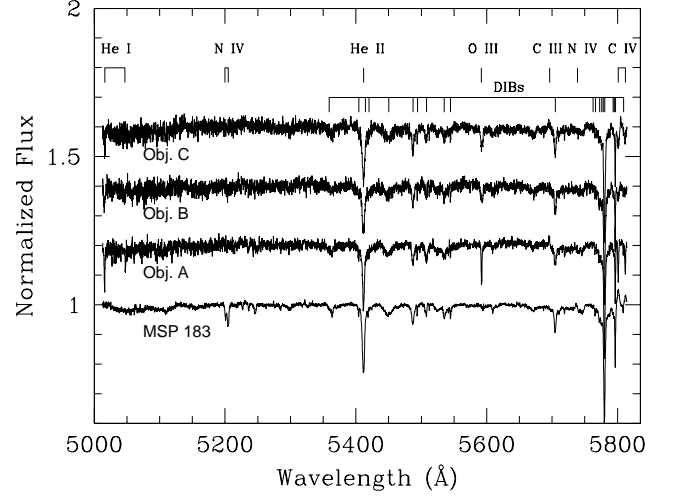


Fig. 12. Mean GIRAFFE-ARGUS spectra of some of the targets discussed in Sect. 3.7.

On the basis of a spectrum in the blue domain, we previously classified MSP 167 as an O6 III star (Paper I). The yellow spectral domain, however, is more consistent with an earlier spectral type. Indeed, our FLAMES-ARGUS spectra reveal no trace of either the He I $\lambda\lambda$ 5015, 5047 or the O III λ 5592 absorption lines. Furthermore, the C IV $\lambda\lambda$ 5801, 5812 lines are clearly not present in absorption. If anything, C IV λ 5801 is in very weak emission ($EW \sim -0.10 \text{ \AA}$). These properties are at odds with an O6 spectral type because the O III λ 5592 and C IV $\lambda\lambda$ 5801, 5812 absorption lines are expected to be strongest for the spectral types O6-O7 (Walborn 1980).

The dominant stellar line is He II λ 5412. We measured its RVs by fitting Gaussian profiles (see Table 6). Whilst there is only little difference between the RVs of the first two spectra, there is a clear shift of more than 100 km s^{-1} between the second and third observation (see Fig. 13). This shift is highly significant since the RV dispersion of the DIBs range between 1.0 and 9.5 km s^{-1} with the lower limit on the dispersion holding for those DIBs which are of similar strength to the He II λ 5412 line.

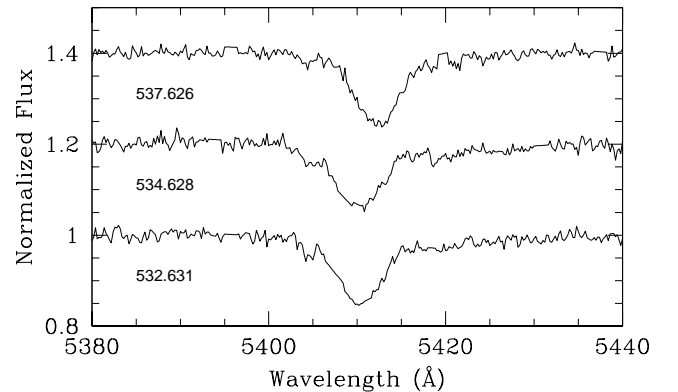


Fig. 13. The He II λ 5412 line in the spectrum of MSP 167 at the three dates of our GIRAFFE-ARGUS observations. The shift between the second and the third observation is clearly seen.

The above results suggest that MSP 167 is actually a short-period (of the order of a week) spectroscopic binary consisting

Table 6. Radial velocities of the binary candidate MSP 167.

| HJD-2454000 | RV (He II λ 5412) (km s ⁻¹) |
|-------------|--|
| 532.631 | -56.2 |
| 534.628 | -64.5 |
| 537.626 | 49.0 |

of an early O-type star (responsible for the strong He II line seen in our ARGUS data) and a companion of later spectral type (which produces the He I absorption lines seen in the blue-band spectra). We note that this star displays a rather large X-ray luminosity of 8.9×10^{32} erg s⁻¹ with some hints of moderate (2σ) variability (see Nazé et al. 2008). These properties, along with the RV variations reported above, suggest that MSP 167 might be an interacting-wind binary system.

For MSP 199 (O3-4 V according to Paper I), the spectra were extracted from a single spaxel of the ARGUS unit. Therefore, the S/N of the individual spectra is rather low. Nevertheless, the combined spectrum agrees with the classification proposed in Paper I. The RVs of the He II λ 5412 display a $1 - \sigma$ dispersion of 5.3 km s^{-1} , whilst the DIBs at λ 5781 and λ 5796 have RV dispersions of 1.9 and 2.4 km s^{-1} , respectively.

Objects A, B, C, D, and E are discussed here for the first time. Our mean spectrum of object A (10:24:02.8 –57:45:30.7, J2000) displays rather strong He II λ 5412, O III λ 5592, and C IV λ 5801, 5812 absorption lines along with fainter absorption lines of He I λ 5016, 5047 and a very weak C III λ 5696 emission line (see Fig. 12). These spectral features are indicative of an O8 spectral classification with an uncertainty of one subtype (Kerton et al. 1999, Walborn 1980). The weakness of the C III emission most likely indicates a main-sequence luminosity class. Our photometric data (Paper I) indicate $V = 15.766 \pm 0.050$ and $B - V = 1.456$ for this star. Assuming $R_V = 3.1$ and adopting the intrinsic colours and absolute magnitudes of an O8 V star (with an uncertainty of one spectral type) from Martins & Plez (2006), we infer a distance modulus 14.83 ± 0.30 corresponding to a spectrophotometric distance of $(9.2 \pm 1.3) \text{ kpc}$. The RVs of the various stellar absorption lines yield $1 - \sigma$ dispersions in the range 1.1 to 6.5 km s^{-1} depending on the strength of the line considered (see Table A.8). The lowest dispersions are found for the He II λ 5412 and C IV λ 5801 lines. The DIBs at λ 5781 and λ 5796 have RV dispersions of 4.1 and 1.3 km s^{-1} , respectively.

For star B (10:24:02.6 –57:45:31.2, J2000), the lines of He I and O III are weak, but clearly present, whilst the C IV absorption lines are also rather weak (see Fig. 12). The strength of the He II λ 5412 line implies a spectral type around O6-7. The RVs of the latter line indicate a $1 - \sigma$ dispersion of 3.8 km s^{-1} , which is slightly larger than the one of the DIB at λ 5781 (1.9 km s^{-1}).

Star C (10:24:02.5 –57:45:32.1, J2000) displays a weak He I λ 5016 line, as well as shallow and broad O III and C IV absorption lines (Fig. 12), and C III λ 5696 is seen in very weak emission. The spectral type is likely about O7 V. The RV dispersion of the He II λ 5412 line amounts to 2.2 km s^{-1} , very much comparable to that of the DIBs at λ 5781 and λ 5797 (2.3 and 1.6 km s^{-1} , respectively).

For stars D and E, the S/N of the spectra is rather low, preventing us from measuring the RVs of the stars from the individual spectra. We instead combined the three observations of each star into a mean spectrum with the goal of performing a spectral classification. Object D (10:24:02.2 –57:45:32.0, J2000) has a stronger He I λ 5016 than He II λ 5412 line. We accordingly classify this star as O9.5 (Kerton et al. 1999). Star E (10:24:02.6 –57:45:32.9, J2000) must be of earlier spectral type, since there is no trace of He I λ 5016, whilst He II λ 5412 is the strongest stellar line in the spectrum. A weak O III λ 5592 line is seen, but no C IV absorption could be detected. We suggest a spectral type near O6-7.

We inspected the *Chandra* image of the cluster’s core (see Nazé et al. 2008) at the positions of objects A – E. Star E corresponds to the southwest lobe of an X-ray source detected between MSP 183 and MSP 199 using the wavedetect algorithm (Naze et al. 2008). Whilst there is no obvious X-ray emission from stars A and C, there might be some X-ray emission associated with stars B and D. However, the latter two objects lie in the wings of the X-ray source associated with MSP 183, and their status as X-ray emitters can therefore not be ascertained with 100% confidence. The lack of X-ray detections for stars A and C is somewhat surprising given their spectral types inferred above. One possibility could be a larger interstellar column density, thus larger absorption of the X-ray emission than for other stars in Westerlund 2 of similar spectral type. However, this explanation needs to be confirmed using additional high angular resolution optical and X-ray observations.

4. Discussion

The various eclipsing binaries that we have studied in this paper yield distance estimates in the range 6.5 to larger than 9 kpc . Whilst the uncertainty in these results is difficult to evaluate, they are in very good agreement with the spectrophotometric distance of $(8.0 \pm 1.4) \text{ kpc}$ inferred in Paper I and clearly consistent with a cluster distance well beyond the 2.8 kpc proposed by Ascenso et al. (2007). We thus conclude that all available observations of the early-type stellar population of Westerlund 2 are most consistent with a distance near 8.0 kpc in agreement with the distance inferred for WR 20a. This result casts serious doubts on a possible association between the γ -ray pulsar PSR J1023.0–5746 (Ackermann et al. 2011) and the Westerlund 2 cluster. If the distance of the pulsar were confirmed at 2.4 kpc , it would then appear more likely that the latter belongs to a foreground stellar population unrelated to the massive stars of Westerlund 2, but possibly related to the PMS stars discussed by Ascenso et al. (2007).

With respect to the multiplicity of early-type stars in Westerlund 2, we stress that among the eleven O-type stars and one WNha star monitored during our campaign, which were not previously known to be binary systems, only one star (MSP 167) was found to probably be a short-period spectroscopic binary system. The other targets did not display RV variations well above the uncertainties estimated from the RV dispersion of the DIBs. To assess the significance of our results, we performed Monte Carlo simulations (Hammersley & Handscomb 1964) of the RVs of a synthetic population of massive binaries. The parent distribution of the binary parameters of the population of massive binaries is unfortunately unknown, but we can make a series of reasonable assumptions. Here, we have followed an approach similar to the one of Sana, Gosset & Evans (2009). Following

Table 7. Results of our Monte Carlo simulations of the detection of massive binaries with the sampling of the GIRAFFE-IFU campaign.

| $m_1 (M_\odot)$ | I | II | III | IV | V |
|-----------------|-------------------|------------------------|-------------------------|-------------------------|--------------------------|
| | $P \leq 5$ (%) | $5 < P \leq 10$ (%) | $10 < P \leq 20$ (%) | $20 < P \leq 50$ (%) | $50 < P \leq 100$ (%) |
| 20 | 0.7 | 0.7 | 1.8 | 19.0 | 56.2 |
| 30 | 0.6 | 0.6 | 1.5 | 15.8 | 50.8 |
| 40 | 0.5 | 0.5 | 1.2 | 13.8 | 46.6 |
| 50 | 0.4 | 0.4 | 1.0 | 12.4 | 43.6 |

Notes. For a given range of orbital periods (expressed in days) and a given value of the primary mass, the table yields the fraction of systems (in % of the population of systems within this interval of periods) that would display $\Delta RV = \max RV(\phi_i) - \min RV(\phi_i) < 20 \text{ km s}^{-1}$, hence escape detection. The total numbers of simulated systems in each of the period intervals are 34198, 25802, 4786, 6315, and 4772 for the intervals I, II, III, IV, and V, respectively.

Table 8. Same as Table 7, but for the sampling of the GIRAFFE-ARGUS data set.

| $m_1 (M_\odot)$ | I | II | III | IV | V |
|-----------------|-------------------|------------------------|-------------------------|-------------------------|--------------------------|
| | $P \leq 5$ (%) | $5 < P \leq 10$ (%) | $10 < P \leq 20$ (%) | $20 < P \leq 50$ (%) | $50 < P \leq 100$ (%) |
| 20 | 3.4 | 1.5 | 4.5 | 29.8 | 70.8 |
| 30 | 2.7 | 1.2 | 3.4 | 26.2 | 64.8 |
| 40 | 2.4 | 1.0 | 2.8 | 23.7 | 60.7 |
| 50 | 2.1 | 0.9 | 2.3 | 21.7 | 57.6 |

the latter authors, we adopt a bi-uniform period distribution in $\log P(\text{days})$ with 60% of the systems in the range $0.3 \leq \log P \leq 1.0$ and the remainder in the range $1.0 < \log P \leq 3.5$. The eccentricity was assumed to be uniformly distributed between 0.0 and 0.9, with all systems with orbital periods shorter than four days assumed to have circular orbits. The longitudes of periastron and true anomaly of the first observation were taken to be uniformly distributed in the $[0, 2\pi]$ interval and we adopted a uniform distribution of the mass ratio $q = m_2/m_1$ between 0.1 and 1.0. The lower limit to the range of mass ratio was chosen because, in practice, it would be very difficult to observe the reflex motion of an O-star for such a low-mass companion. Finally, the orbital inclination was taken to be uniform in $\cos i$ between -1.0 and 1.0 .

For four values of the primary mass (20, 30, 40, and $50 M_\odot$, spanning roughly the range of spectral types for which some observational information on multiplicity is available) we simulated a population of 100000 binary systems. For each system, we evaluated the maximum RV difference $\Delta RV = \max RV(\phi_i) - \min RV(\phi_i)$ that would be measured on this system adopting the same temporal sampling as used for our five GIRAFFE-IFU observations⁸. The results are summarized in Table 7 and illustrated in Fig. 14. If we assumed that a velocity difference of 20 km s^{-1} were required to claim the detection of a binary system⁹, we would find that fewer than 1.5% of the systems containing an O4-6 V primary star (i.e. with a mass higher than $30 M_\odot$), and having an orbital period of shorter than 20 days, would escape detection. Only the lowest inclination systems would indeed remain undetected. For longer orbital periods, the fraction of systems that would be missed by our GIRAFFE-IFU campaign steeply increases: for systems with orbital periods between 50

and 100 days, about half of the binaries would escape detection. However, owing to the assumed preponderance of short orbital-period systems, the total fraction of systems with a period up to 100 days that we would miss, would still remain quite low (4 - 5%).

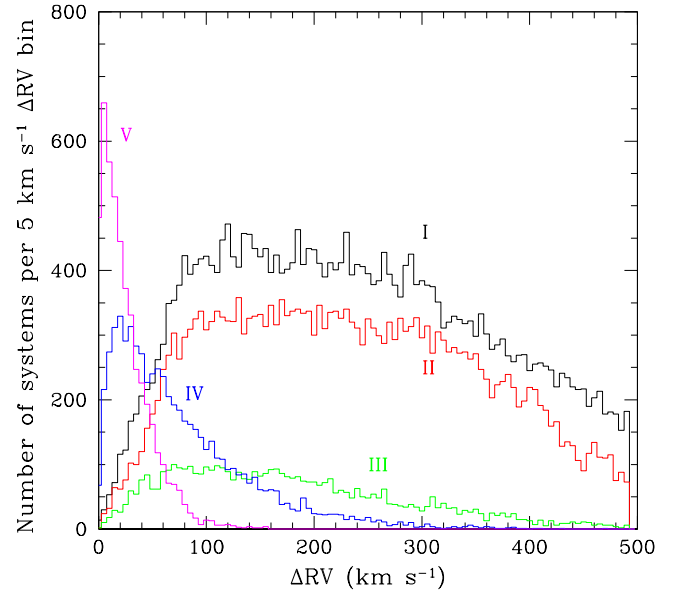


Fig. 14. Distribution of $\Delta RV = \max RV(\phi_i) - \min RV(\phi_i)$ for a population of massive binaries subdivided into various intervals of orbital period (labelled I, II, III, IV, and V, see Table 7). In this simulation, m_1 was assumed to be equal to $40 M_\odot$.

⁸ Our UVES campaign includes eight pointings. However, the three additional pointings only yield a marginal improvement in the detection efficiency compared to the results discussed here.

⁹ None of the presumably single O-star objects discussed in this paper displays a velocity excursion of more than 20 km s^{-1} in the most stable photospheric lines.

The same approach was applied to the sampling of the three GIRAFFE-ARGUS observations (see Table 8). As could be expected, this sampling is less efficient for the detection of binary systems. Nonetheless, for primary masses in the range 20

- $50 M_{\odot}$ (corresponding roughly to O4-8 V spectral types), our simulations indicate that a binary system with a period shorter than 100 days has less than 10% a probability of escaping detection because $\Delta RV < 20 \text{ km s}^{-1}$.

Therefore, we can safely conclude that the probability that either of the stars MSP 18, 171, 182, 183, 199, or 203 or stars A, B or C is a short-period binary system must be very low. If these stars are indeed binaries, they must have either a very low orbital inclination, a very long orbital period, or both. The situation is a bit less clear for WR 20b for which we do observe a maximum RV difference above our threshold. However, as stated above, this result must be considered with caution, since all the strong emission lines display line-profile variability that could easily be responsible for the observed shifts in RV, whilst the weaker emission lines are found to be more constant, in terms of both morphology and RV.

The lack of a clear binary signature for WR 20b and MSP 18 is especially surprising given their high X-ray luminosities (see Nazé et al. 2008). These stars have $\log L_X/L_{\text{bol}} = -6.16$ and -5.92 for WR 20b and MSP 18 respectively, which are significantly larger than for typical single O-type stars and are as large as, or even larger than the value of the interacting wind system WR 20a ($\log L_X/L_{\text{bol}} = -6.15$, Nazé et al. 2008). Since colliding winds have the potential to generate copious amounts of X-rays (e.g. Stevens, Blondin & Pollock 1992), high X-ray luminosities have traditionally been associated with interacting winds in a binary system. However, recent investigations have revealed that interacting wind binaries are not always X-ray overluminous (e.g. Sana et al. 2006, Nazé 2009), at least as far as O-star binaries are concerned, and that alternative mechanisms such as magnetically confined winds can also produce high X-ray luminosities (e.g. Babel & Montmerle 1997). This implies that colliding winds, and by extension binarity, are neither a necessary nor a sufficient condition for high X-ray luminosities. Nonetheless, the high X-ray luminosities of MSP 18 and WR 20b make these objects very interesting targets that deserve additional monitoring, in terms of both optical spectroscopy, to search for long-term RV variations, and spectropolarimetry, to search for evidence of magnetic fields or non-spherical winds.

If these stars were indeed single, then their bolometric luminosities would suggest masses in the range $50 - 70 M_{\odot}$, provided that the main-sequence mass-luminosity relation holds for all these objects. It has been argued that the most massive stars might actually display an O If* supergiant or hydrogen-rich WN type spectrum rather than a normal O-type spectrum (Crowther et al. 2010). While it is true that the highest ever dynamical masses have been measured for WNLha stars so far (Rauw et al. 2004, Niemela et al. 2008, Schnurr et al. 2008, 2009), it is puzzling that the WN6ha star WR 20a and the O5 V-III stars MSP 18 and MSP 203 have very similar bolometric luminosities and hence masses. In this context, it is worth recalling the result of De Becker et al. (2006), and Taylor et al. (2011). De Becker et al. (2006) derived a minimum dynamical mass of $(150 \pm 50) M_{\odot}$ for the O5.5 III(f) primary of HD 15558, whilst Taylor et al. (2011) inferred minimum dynamical masses of $M_1 \sin^3 i = (78 \pm 8)$ and $M_2 \sin^3 i = (66 \pm 7) M_{\odot}$ for the O6.5 Iafc + O6 Iaf binary system R 139. After all, these results suggest that some apparently ‘normal’ O-type stars could actually span a range of masses quite similar to that of WNLha stars.

Although we have monitored only a limited sample of O-type stars in Westerlund 2, our results suggest that unlike some other very young open clusters, such as NGC 6231 (Sana

et al. 2008) or Trumpler 16 (Rauw et al. 2009), Westerlund 2 is most probably not very rich in short period early-type binaries. With only five binaries or binary candidates (including WR 20a and the three eclipsing OB systems, two of which host B-type primaries) out of the fifteen stars investigated, our results are more reminiscent of the situation encountered in IC 1805 (De Becker et al. 2006, Hillwig et al. 2006) and NGC 2244 (Mahy et al. 2009). If we restrict ourselves to O-type stars, the fraction of confirmed or candidate binaries drops to about 15% (2 out of 13). Given the higher stellar density in the core of NGC 6231 than in both NGC 2244 and IC 1805, Mahy et al. (2009) tentatively suggested that the number of short-period early-type binary systems might be correlated with the cluster density. However, with its rather large cluster density and low fraction of short-period massive binaries, Westerlund 2 apparently does not follow such a correlation. A more extensive multi-epoch monitoring of the other O-type stars in Westerlund 2 would be highly valuable in helping us to establish the true binary fraction and investigate the question of a positional dependence of the occurrence of early-type binaries inside the cluster.

The vast majority of the O-type stars in our sample indicate a RV of about $20 - 30 \text{ km s}^{-1}$ as measured using the He II $\lambda 5412$ line. There are two exceptions to this rule, MSP 18 and MSP 171, which are instead found to have negative RVs for this line of $(-1.1 \pm 1.8) \text{ km s}^{-1}$ and $(-9.3 \pm 6.9) \text{ km s}^{-1}$ for MSP 18 and MSP 171, respectively. Both stars are located outside the dense core of the cluster, at angular separations of 55 and 23'' for MSP 18 and MSP 171, respectively. Therefore, these objects might have been ejected from the cluster core by means of dynamical interactions (see e.g. Leonard & Duncan 1990) or be long-period eccentric binaries observed near their quadrature phase. Alternatively, these stars could have strong stellar winds and their He II $\lambda 5412$ line might then be blue-shifted owing to contamination by wind effects. This latter explanation is however at odds with the absence of other strong wind signatures in their spectra (see the discussion of their spectra in the present work and in Paper I). We therefore conclude that these stars are runaway candidates or possible long-period eccentric binaries observed near a phase of maximum RV excursion.

Finally, we have obtained the very first spectra of mid- to late-type O stars (MSP 223, A, B, C, D, and E) in Westerlund 2. Together with the results of Paper I, the number of known O-type stars in the cluster is now 18. However, given the apparent magnitude of these mid to late-type O stars, only the tip of the iceberg has yet been identified. Our current census of O-type stars in Westerlund 2 is obviously still far from being complete and requires additional high spatial-resolution spectroscopic investigations. Indeed, from the relationship between the mass of the most massive star and the mass of its parental cluster (Weidner et al. 2010), we estimate that if $m_{\text{max}} = 85 M_{\odot}$, Westerlund 2 likely contains about $3000 M_{\odot}$ in the form of stars. Adopting the IMF of Weidner et al. (2010), we then estimate that the number of stars more massive than $20 M_{\odot}$ should be around 45. Therefore, more than half of the O-type stars of Westerlund 2 would still await discovery!

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Appendix A: Tables of individual RV measurements

The tables below provide the heliocentric RVs measured using various lines of stars that do not exhibit any clear binary signature.

Table A.1. Radial velocities of WR 20b as measured for the various emission lines in the spectrum.

| Date HJD-2454000 | Radial velocity | | | | | | | | |
|---------------------|------------------------------------|---|--|---|---|--|---|-------------------------------------|---|
| | H β (km s ⁻¹) | N v λ 4945 (km s ⁻¹) | N iv λ 5204 (km s ⁻¹) | He ii λ 5412 (km s ⁻¹) | N v λ 5737 (km s ⁻¹) | N iv λ 6383 (km s ⁻¹) | He ii λ 6406 (km s ⁻¹) | H α (km s ⁻¹) | He ii λ 6683 (km s ⁻¹) |
| 532.587 | 132.2 | -22.0 | 57.8 | 108.0 | -52.9 | 21.0 | -15.8 | 71.1 | 33.9 |
| 532.631 | 125.2 | -23.2 | 33.6 | 105.9 | -34.7 | -1.2 | -27.4 | 71.5 | 41.2 |
| 534.581 | 101.9 | -49.5 | 36.6 | 87.2 | -58.8 | 28.9 | -30.2 | 63.3 | 3.1 |
| 534.628 | 91.2 | -20.9 | 29.0 | 78.7 | -48.3 | 7.4 | -24.5 | 68.9 | 11.7 |
| 536.624 | 84.3 | -45.0 | 49.0 | 95.0 | -72.9 | 46.8 | -15.3 | 11.9 | 44.8 |
| 537.626 | 125.7 | -39.4 | 55.0 | 105.1 | -38.5 | -14.2 | 0.6 | 43.5 | 46.0 |
| 538.719 | 132.4 | -42.2 | 74.3 | 123.9 | -58.7 | 13.5 | 7.4 | 100.5 | 66.6 |
| 539.622 | 132.1 | -36.4 | 64.8 | 125.7 | -38.9 | 46.4 | 12.8 | 125.8 | 65.0 |

Table A.2. Radial velocities of MSP 18 measured using various absorption lines and the weak C iii λ 5696 emission line.

| Date HJD-2454000 | Radial velocity | | | | |
|---------------------|------------------------------------|---|---|-------------------------------------|---|
| | H β (km s ⁻¹) | He ii λ 5412 (km s ⁻¹) | C iii λ 5696 (km s ⁻¹) | H α (km s ⁻¹) | He ii λ 6683 (km s ⁻¹) |
| 532.587 | -41.7 | -3.9 | -16.8 | -122.1 | 4.2 |
| 532.631 | -46.1 | 1.7 | -0.9 | -123.8 | 9.1 |
| 534.581 | -45.5 | -2.1 | -17.0 | -122.3 | 9.9 |
| 534.628 | -33.5 | -2.0 | -17.3 | -118.4 | 14.9 |
| 536.624 | -42.8 | -2.0 | 1.0 | -122.7 | 6.3 |
| 537.626 | -34.3 | 0.6 | -8.3 | -120.3 | 19.2 |
| 538.719 | -50.1 | -0.5 | -6.8 | -122.3 | 6.4 |
| 539.622 | -44.6 | -0.3 | 7.2 | -122.5 | 0.8 |

Table A.3. Radial velocities of MSP 171 measured using various absorption lines.

| Date HJD-2454000 | Radial velocity | | | |
|---------------------|---|---|--|--|
| | He ii λ 5412 (km s ⁻¹) | O iii λ 5592 (km s ⁻¹) | C iv λ 5801 (km s ⁻¹) | C iv λ 5812 (km s ⁻¹) |
| 532.587 | -0.1 | -10.5 | 10.3 | -13.5 |
| 534.581 | -14.4 | -16.9 | -7.8 | -22.1 |
| 536.624 | -17.5 | -2.6 | -5.0 | -17.4 |
| 538.719 | -8.6 | -23.1 | 5.5 | -21.7 |
| 539.622 | -5.8 | -20.5 | 5.5 | -18.3 |

Table A.4. Same as Table A.3 but for MSP 182.

| Date HJD-2454000 | Radial velocity | | | |
|---------------------|---|---|--|--|
| | He ii λ 5412 (km s ⁻¹) | O iii λ 5592 (km s ⁻¹) | C iv λ 5801 (km s ⁻¹) | C iv λ 5812 (km s ⁻¹) |
| 532.587 | 34.5 | 28.9 | 36.8 | 4.1 |
| 534.581 | 27.7 | — | 36.5 | 8.0 |
| 536.624 | 33.3 | 30.4 | 43.2 | 1.3 |
| 538.719 | 29.8 | 29.8 | 38.0 | 16.7 |
| 539.622 | 27.0 | 35.6 | 27.5 | -2.8 |

Table A.5. Same as Table A.3 but for MSP 203 + MSP 444.

| Date HJD-2454000 | Radial velocity | | | |
|---------------------|---|---|--|--|
| | He ii λ 5412 (km s ⁻¹) | O iii λ 5592 (km s ⁻¹) | C iv λ 5801 (km s ⁻¹) | C iv λ 5812 (km s ⁻¹) |
| 532.587 | 28.4 | 47.7 | 27.3 | -0.8 |
| 534.581 | 24.7 | 38.8 | 25.6 | 5.5 |
| 536.624 | 23.3 | 33.7 | 21.1 | 7.4 |
| 538.719 | 21.5 | 39.4 | 17.4 | 0.0 |
| 539.622 | 23.9 | 52.6 | 19.8 | 0.6 |

Table A.6. Radial velocities of MSP 183 measured using various absorption and emission (E) lines.

| Date HJD-2454000 | Radial velocity | | |
|---------------------|---|--|--|
| | He II λ 5412 (km s ⁻¹) | C IV λ 5801 (E) (km s ⁻¹) | C IV λ 5812 (E) (km s ⁻¹) |
| 532.631 | 23.0 | 38.0 | 53.9 |
| 534.628 | 25.3 | 31.3 | 55.5 |
| 537.626 | 23.5 | 23.4 | 49.1 |

Table A.7. Radial velocities of MSP 199, Star B and C measured using the He II λ 5412 line.

| Date HJD-2454000 | Radial velocity | | |
|---------------------|----------------------------------|---------------------------------|---------------------------------|
| | MSP 199 (km s ⁻¹) | Star B (km s ⁻¹) | Star C (km s ⁻¹) |
| 532.631 | 29.1 | 18.0 | 28.0 |
| 534.628 | 36.4 | 13.6 | 29.7 |
| 537.626 | 26.2 | 21.3 | 32.3 |

Table A.8. Same as Table A.3 but for Object A.

| Date HJD-2454000 | Radial velocity | | | | |
|---------------------|--|---|---|--|--|
| | He I λ 5016 (km s ⁻¹) | He II λ 5412 (km s ⁻¹) | O III λ 5592 (km s ⁻¹) | C IV λ 5801 (km s ⁻¹) | C IV λ 5812 (km s ⁻¹) |
| 532.631 | 20.6 | 26.3 | 17.1 | 23.0 | 25.6 |
| 534.628 | 8.8 | 28.1 | 22.5 | 22.0 | 26.6 |
| 537.626 | 19.4 | 28.5 | 16.5 | 18.7 | 19.2 |